Imaging the Solar Wind in 3D with the PUNCH Constellation of Small Satellites

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ABSTRACT: The Polarimeter to UNify the Corona and Heliosphere (PUNCH) is the latest Small Explorer mission being developed for NASA. Beginning in 2023, PUNCH will use polarized visible-light imaging to understand how the Sun's corona becomes the solar wind that fills our solar system. The space segment of the mission is a constellation of four small satellites (~50kg each) that work together to form a "virtual" instrument" with a 90°-wide field of view centered on the Sun. The physical instruments comprise a coronagraph ("Narrow Field Imager", NFI) and three heliospheric imagers ("Wide Field Imager", WFI) that together sweep out the entire field of view. The instruments are sensitive to polarization, to enable imaging bright features in 3D using the physics of Thomson scattering. PUNCH is being built by a partnership of Southwest Research Institute, the U.S. Naval Research Laboratory, and the Rutherford Appleton Laboratory.









Mysteries of the Corona-Solar Wind Connection









PUNCH Mission Profile

Mission Structure 1+3 constellation of smallsats **Orbit** 600km 6am/6pm Sun-synch LEO Launch 2023 Q1 **Duration** 2 years (+90 days commissioning) **Spacecraft type** 3-axis stable; propulsion for orbit trim **Development Strategy** Spacecraft are interchangeable; each carries one PUNCH camera

Detailed structure shapes the outer corona and young solar wind:

Recent results from STEREO/ COR2 show that the transition to the solar wind is highly structured at all optically resolvable scales. This contrasts with current solar wind models that treat the flow as smooth. With 10x higher sensitivity than COR2, PUNCH reveals how this structure imprints on the solar wind.

Evolving

Solar Wind

Flow Rates

(1A)

Intermittent structure traces solar wind flow & causes variability at Earth:

Flow in the outer corona is highly variable. Intermittent ejecta may comprise up to 100% of the outflowing plasma. PUNCH is uniquely able to routinely track all of these features in 3D, to trace the origin of the solar wind, identify late-stage acceleration, and determine the role of ejecta in the solar wind as a whole.

The Alfven zone and onset of turbulence mark the elusive base of the heliosphere:

The transition from structured coronal flow to isotropic turbulent flow marks the outer boundary of the corona and the inner boundary of the solar wind. The newly-discovered Alfven zone determines the Sun's open flux budget and shapes the solar wind. PUNCH explores the otherwise inaccessible cross-scale physics of this "discovery space" and the plasma's journey from corona to solar wind.



Polarized imaging reveals how CMEs and solar wind structures move and evolve in 3D

CMEs and other ejecta are fundamentally 3D objects. Deep-field polarized imaging with PÚNCH reveals the 3D details of these structures for the first time, revealing: interior evolution; flux rope chirality (tied to leading-edge Bz); 3D propagation, acceleration and possible deflection; and solar wind interaction of CMEs and ejecta.

PUNCH 3D results have immediate and obvious application (R2O) to space weather forecasting.

Full-field imaging, 4 min. cadence **Concept of Operations** (No targeting, no campaigns), Synchronous across constellation

Field of View All position angles, $1.25^{\circ}-45^{\circ}$ from Sun (90° dia.). NFI: $1.25^{\circ}-8.0^{\circ}$; WFI: $5^{\circ}-45^{\circ}$

Wavelength Range White light (450-750 nm)

Data Products Full field image mosaics (B and pB) Coronal close-up images (B and pB) Background-subtracted "K" images (B/pB) Solar wind flow maps

Data Distribution Via VSO and SDAC; open data policy

PUNCH Instruments

Narrow Field Imager (NFI): Externally occulted coronagraph • Single stage design is simplified compared to STEREO/COR2 • FOV is 1.25° to 8° from Sun • Polarizing filter wheel behind focusing optics • Made at NRL (camera: RAL)

Camera

the sky plane (non-impacting)

Actual Trajectory

CME Distance

from Sun (AU)

ii)Degree of polarizatio

pB/B





FOV and wind maps

One "virtual instrument", 4 cameras

20° 30° 40

brightness varies by over 1,000x. This dynamic range requires two instrument types: a narrow-field imager (NFI) coronagraph similar to STEREO/COR2 or LASCO/C3, and a

in the heliosphere.

Sensitivity and Exposure Time

Current heliospheric imagers are not limited by photon

In the corona, COR2 is photon limited, but its duty cycle is

0.6% in the synoptic observing program. PUNCH uses a

60% shutter duty cycle (100x more exposure), for 10x

ming diff : orig data

A Composite Field of View

Noise=2.4x10⁴⁷ B₀ (1° ay.) Noise=2.5x10⁴⁷ B₄ (1° ay

Running diff.: short ex-

reduction in shot noise compared to COR2.

3D imaging via polarization

Global images for context & completeness

Understanding the young solar wind requires global images of the inner heliosphere. The PUNCH FOV extends from the mid corona to 45° from the Sun at all solar position angles, including high latitudes that are now completely unobserved.





PUNCH maps solar wind flow and acceleration at all azimuths, out to 100 Rs, every few hours. Global solar wind measurement is critical for understanding late-stage wind acceleration and the large- and cross-scale interplay between coronal physics and solar wind conditions.



Observing Geometry: 1 NFI +3 WFIs PUNCH is a constellation of four smallsats distributed around Sun-synchronous twilight orbit. In this geometry, the Sun is always near the horizon. NFI looks directly at the Sun (blocked by its occulting disk), and WFI looks between the Sun and the zenith. A spacecraft roll program maintains

To Sun

Multiple Instruments

Across the PUNCH FOV, the corona and solar wind |

wide-field imager (WFI) similar to STEREO/HI. The

instruments are matched in wavelength and are operated synchronously, to yield a single well-defined field of view.





Polarization for precise 3-D imaging PUNCH measures the location and shape of ejecta in 3D,

using deep-field imaging and the polarization physics of Thomson scattering. This image from STEREO/COR2 demonstrates the technique, but 3x-10x lower noise levels are required for routine polarimetric location of blobs, puffs, and features within CME cores.



CME Trajectory

Extended objects such as CMEs are compact along the line

of sight near the bright leading edge, due to the Mean Value Theorem. PUNCH locates the leading edge of CMEs

directly in 3D by analytic inversion, and characterizes the

3D structure and trajectory using forwad modeling and fitting. PUNCH trajectory analysis is sensitive enough to

detect expected effects such as acceleration, deceleration,

and deflection by the ambient solar wind.

Stereoscopic imaging from two vantage points is limited in the extended, optically thin context of solar wind imaging. In particular, the "leading edge" of a CME or CIR is different from different vantage points, spoiling the stereoscopic inversion. PUNCH overcomes this difficulty with singlevantage 3D measurement of CMEs, CIRs, shocks, and small ejecta.

Overcoming the limits of stereoscopy



CME Chirality

With sufficient signal-to-noise ratio (SNR), it is possible to

forward model of two similar flux rope CMEs with opposite

chirality launched from a model Sun at 45° from the sky

plane. Measuring chirality is a key step toward prediction of

Bz -- a critical measure of CME geoeffectiveness. PUNCH

measure the chirality of a CME, using density tracers to

follow the magnetic field. This demonstration uses a

will demonstrate this technique.

How PUNCH Reveals 3D Structure

